

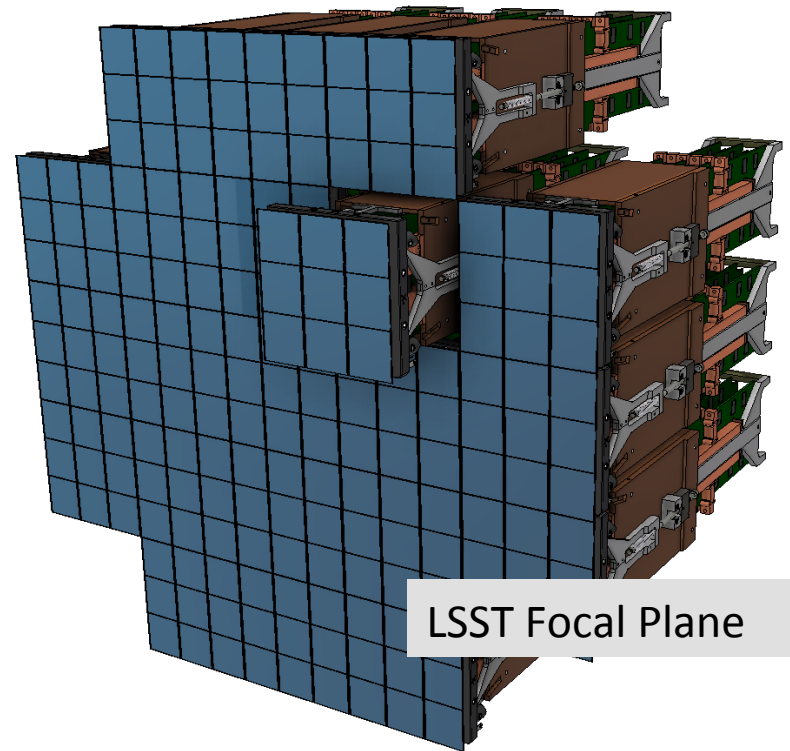
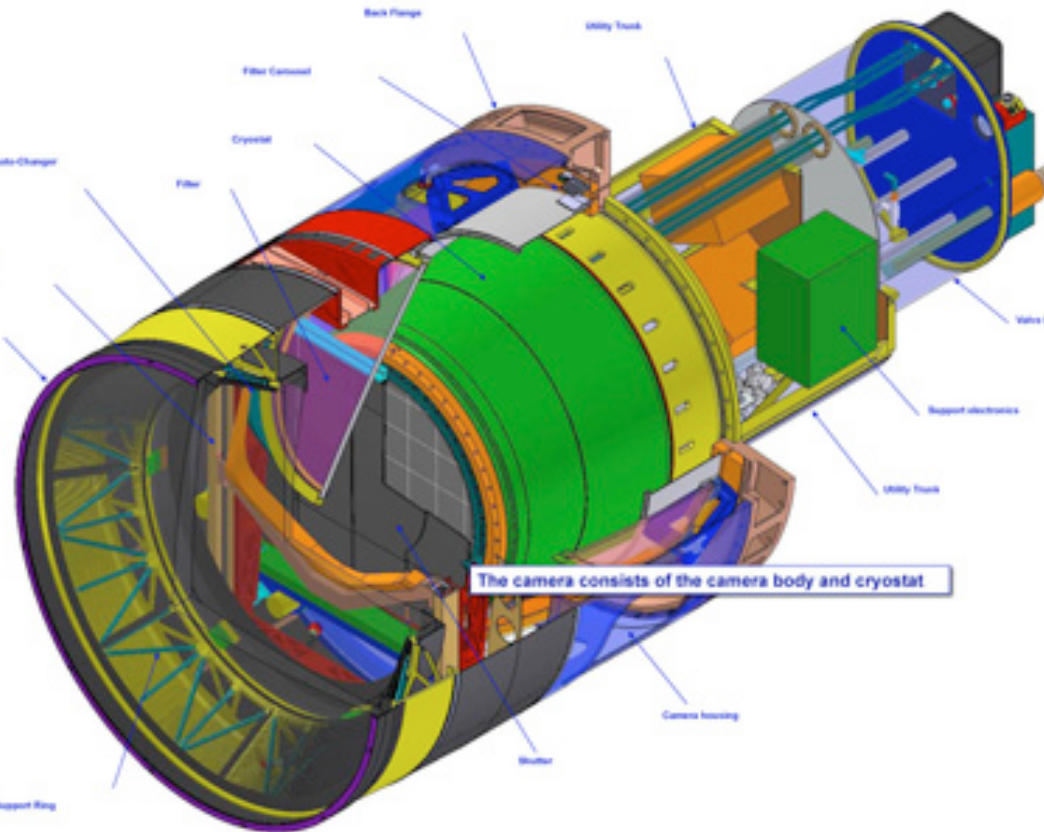
CTE and other sensor effects in thick, fully depleted CCDs

Andrei Nomerotski
Brookhaven National Laboratory

SDW2017, Baltimore
28 September 2017

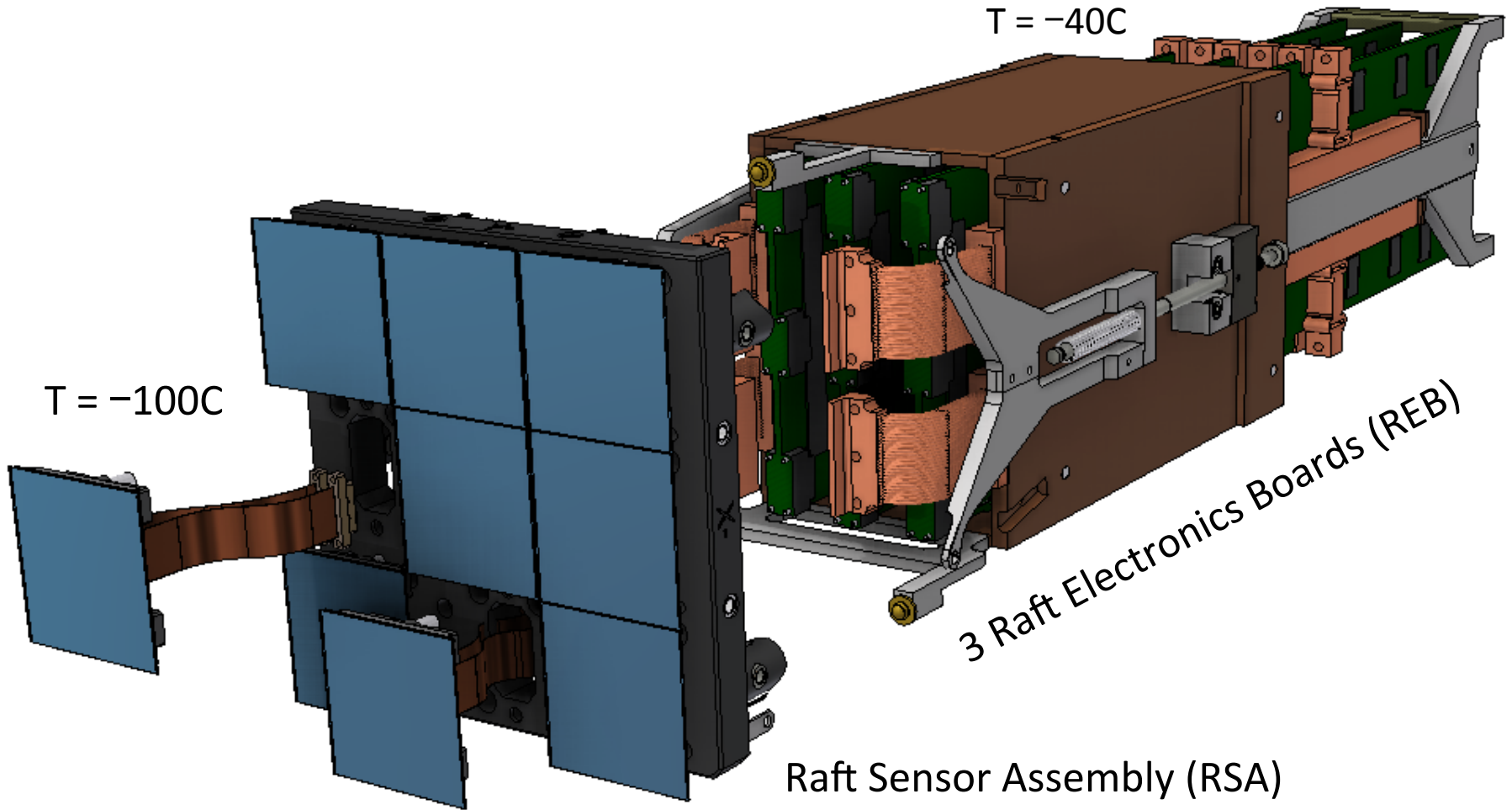
LSST Camera

The Camera Design Overview



- 64cm diameter \rightarrow 3.5°
- 189 CCDs \rightarrow 3.1 Gigapixels
- 21 "rafts" with integrated electronics for nine CCDs

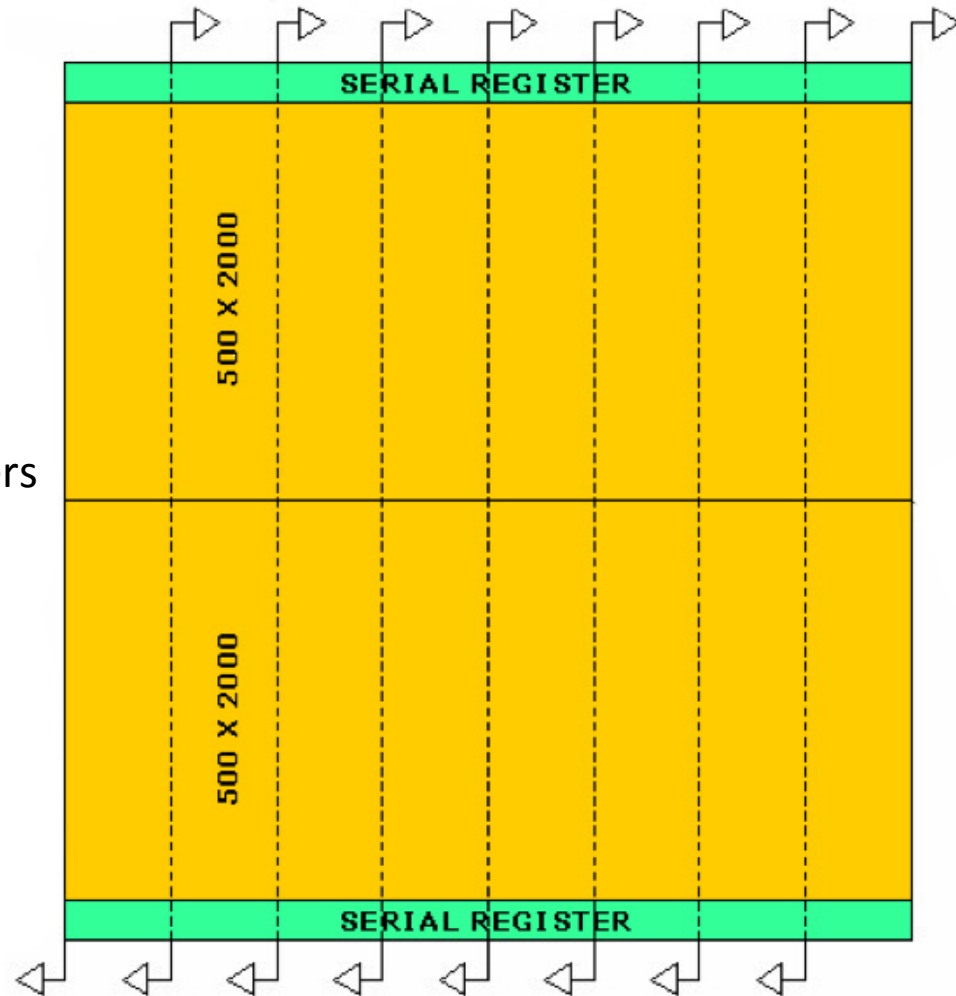
Raft Tower Module



See Paul O'Connor's talk

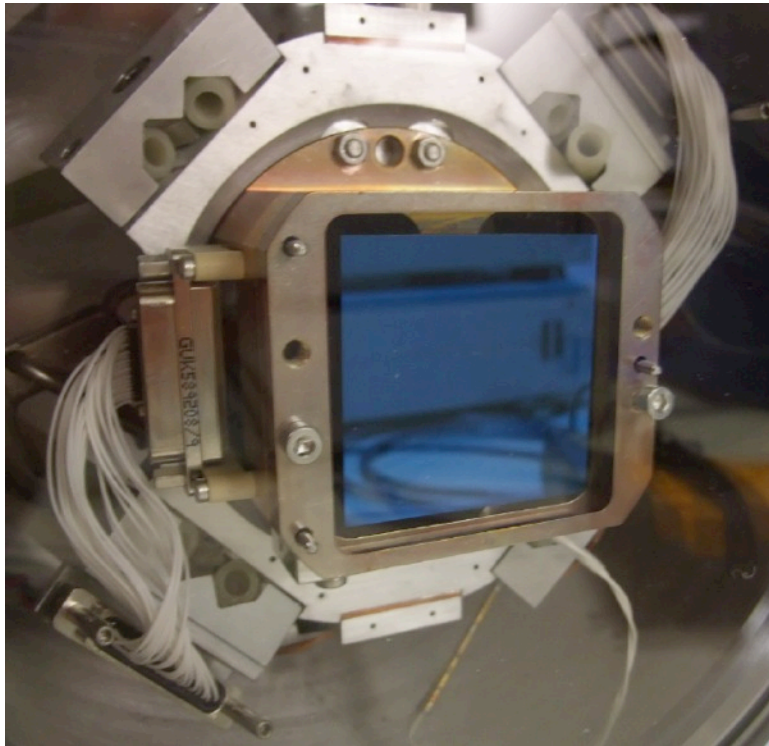
LSST CCD Sensor

- 4k x 4k = 16 Mpixels
- 10x10 microns pixel
- Noise spec 8 e-, based on anticipated sky noise; limits pixel rate
- Pixel read rate is 550 Kpix/s
- 2 second readout time spec → 16 amplifiers
- Si thickness 100 micron →
Enhanced infrared response



LSST CCDs

CCD250



STA3800



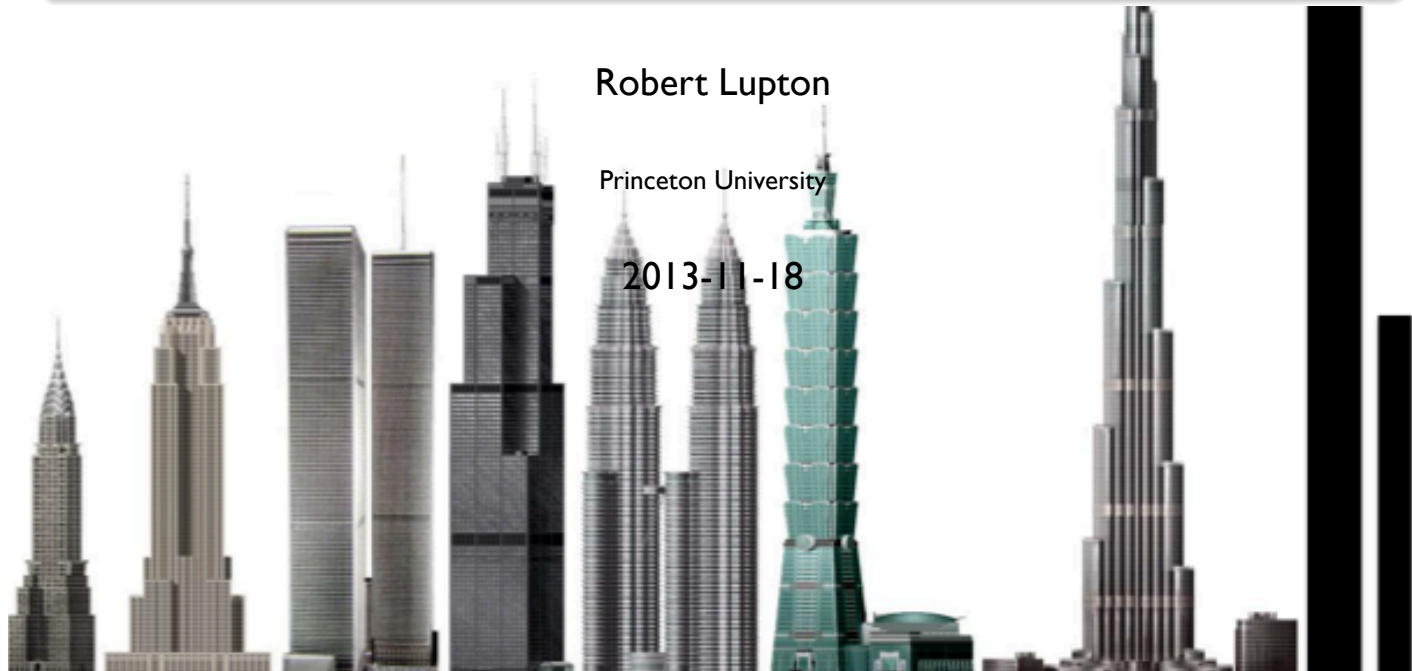
On to more subtle things

LSST pixel : 10 x 10 x 100 micron³

→ Pixels are skyscrapers

1930	1931	1972/73	1974	1997	2004	2010	2013	2022
Chrysler New York 1046 ft 77 Stories	Empire State New York 1250 ft 102 Stories	World Trade Center New York 1368 ft 110 Stories	Sears Tower Chicago 1450 ft 110 Stories	Petronas Towers Kuala Lumpur 1483 ft 88 Stories	Taipei 101 Taipei 1,671 ft 101 Stories	Burj Khalifa Dubai 2717 ft 162 Stories	HSC	LSST

Consequences of thick CCDs on Image Processing

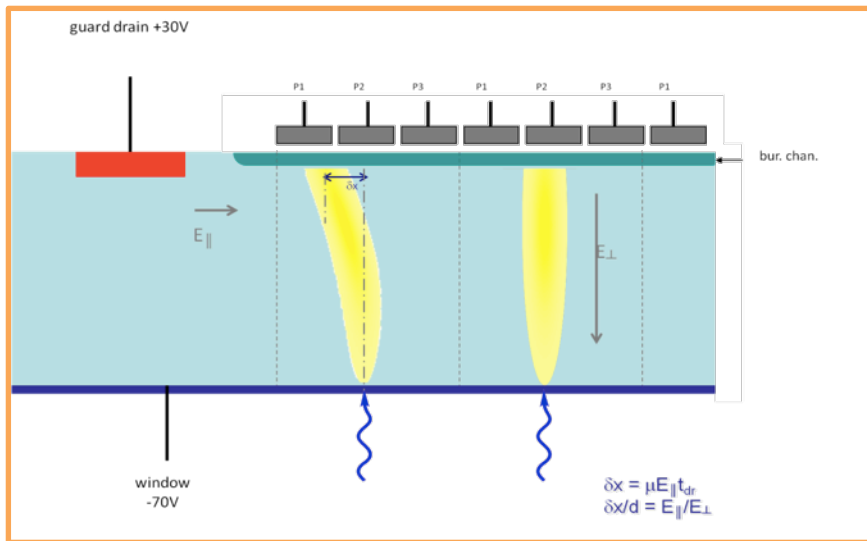


What's the problem with thick CCDs?

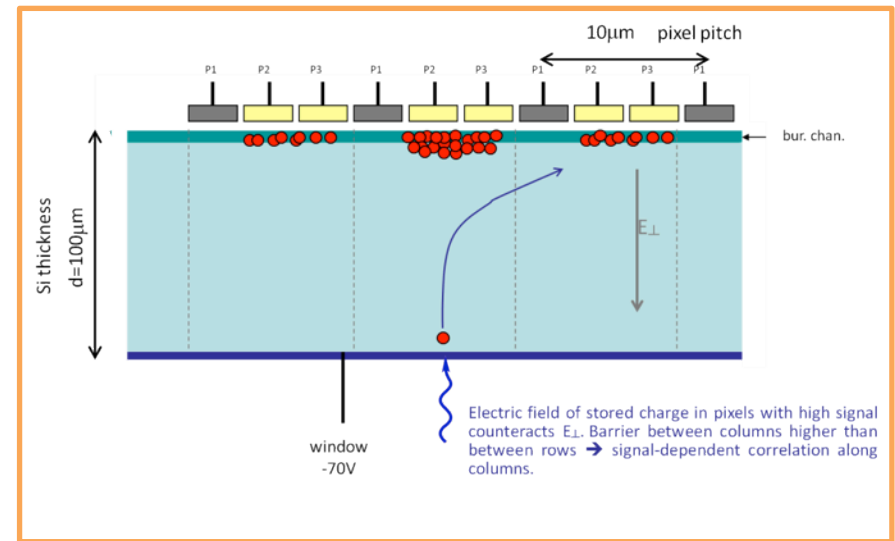
Fully depleted CCD have a non-trivial electrostatics which lead to astrometric biases and PSF distortions

- Difficult to disentangle from photometric effects

Static : edge effects, tree-rings



“Dynamic” : brighter-fatter effect



LSST Dark Energy Science Collaboration – DESC

Sensor Anomalies Working Group

- Main focus on CCD signatures, important for precision astrometry, photometry and shear measurement in LSST, WL is one of strongest motivations

Workshops “Precision Astronomy in Fully Depleted CCDs” at BNL



The Third Workshop
Precision Astronomy
with Fully Depleted CCDs

December 1–2, 2016
BNL Physics Department
Large Seminar Room

The Third Workshop on Precision Astronomy with Fully Depleted CCDs will focus on topics of making precision astronomical measurements of flux (photometry), position (astrometry) and shapes (shear) with thick, fully depleted CCD detectors. Impact of sensor related effects on the Dark Energy science will be addressed through presentations and discussions.

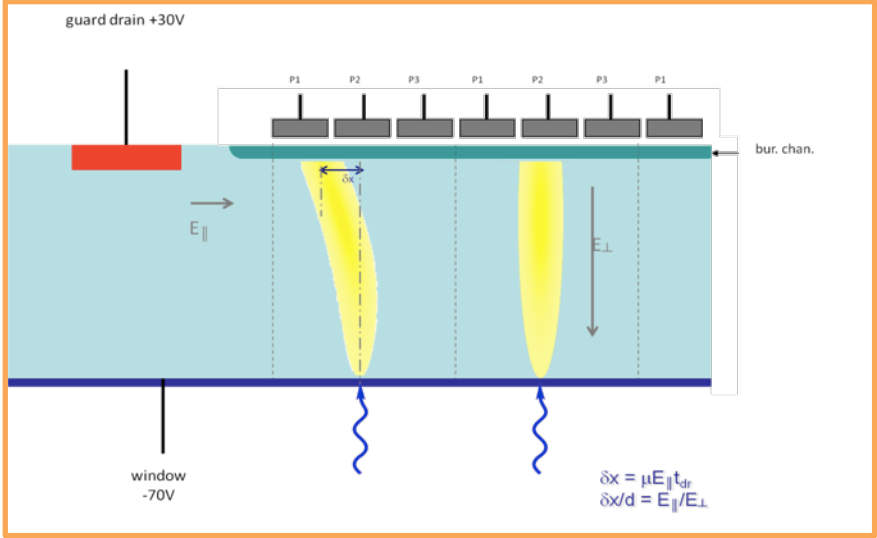
<http://www.bnl.gov/paccd2016>

Organizing Committee
Pierre Astier (LPNHE)
Chris Bebek (LBNL)
Gary Bernstein (Penn)
Juan Estrada (Fermilab)
Mike Jarvis (Penn)
Robert Lupton (Princeton)
Eugene Magnier (Hawaii)
Satoshi Miyazaki (NAOJ)
Andrei Nomerotski (BNL), Chair
Paul O'Connor (BNL)
John Peterson (Purdue)

3 workshops: in 2013, 2014 and 2016

gathering of experts in CCDs and reduction algorithms, peer-reviewed proceedings in JINST

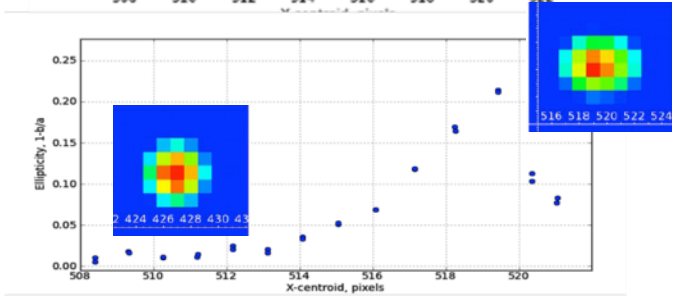
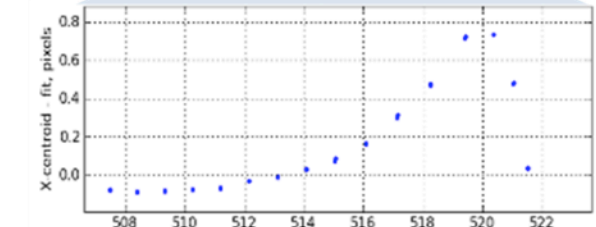
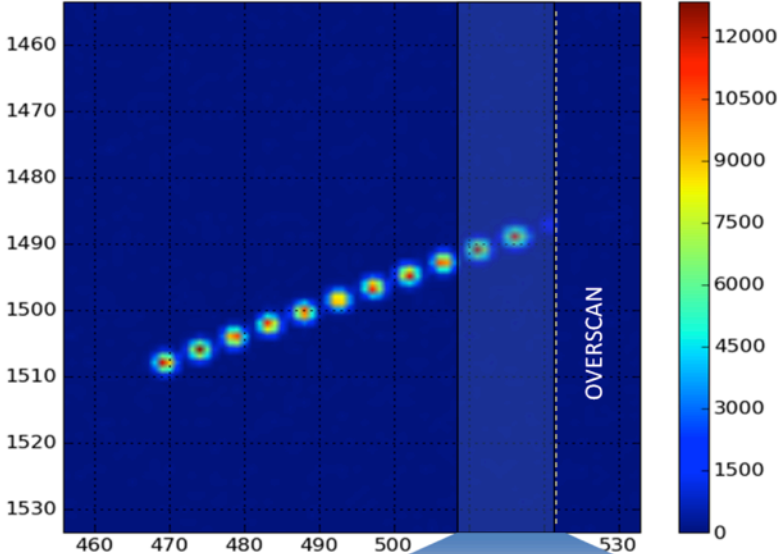
Astrometric Distortions in Thick CCDs



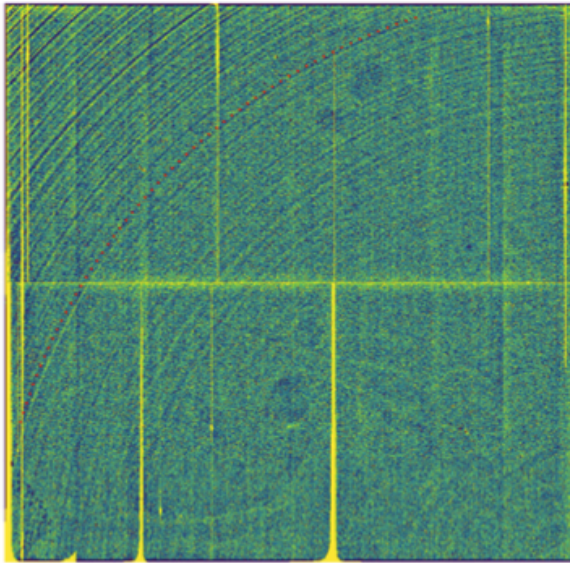
P O'Connor 2014 JINST 9 C03033

Distortions on the edge:

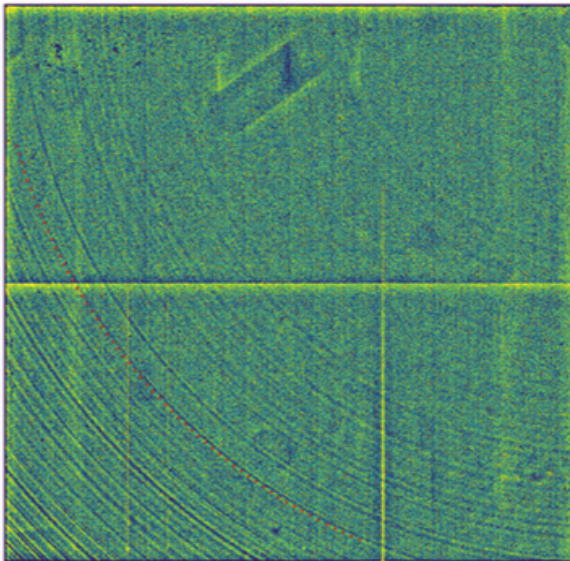
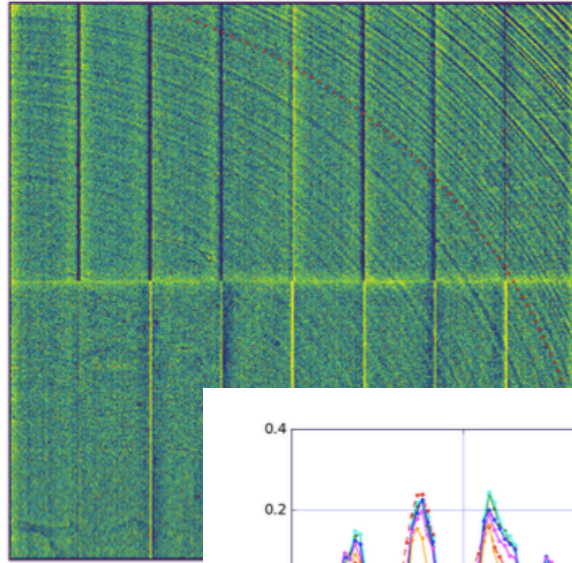
- Astrometric bias: up to 50%
- Ellipticity: up to 20%



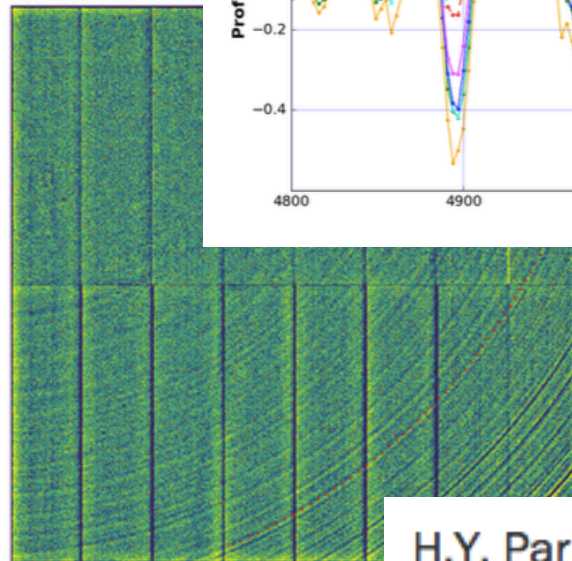
Tree Rings in ITL sensors



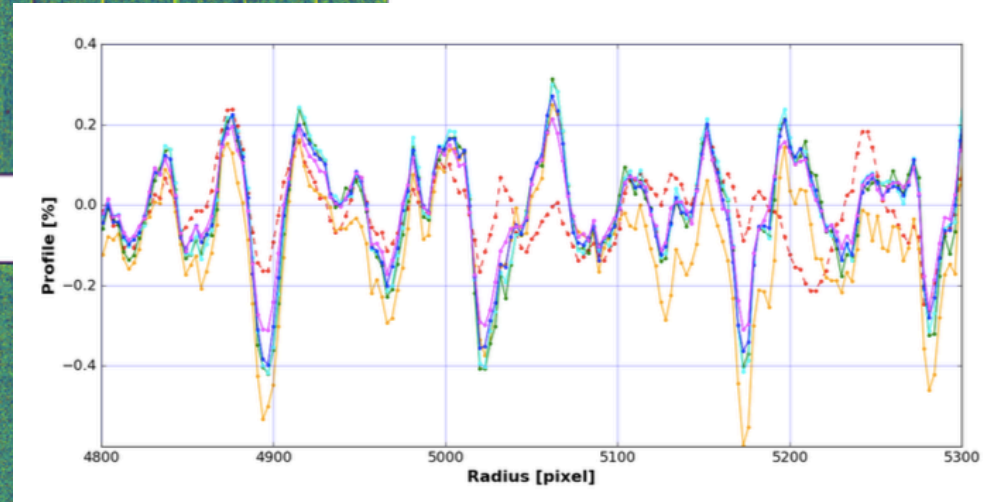
(a)



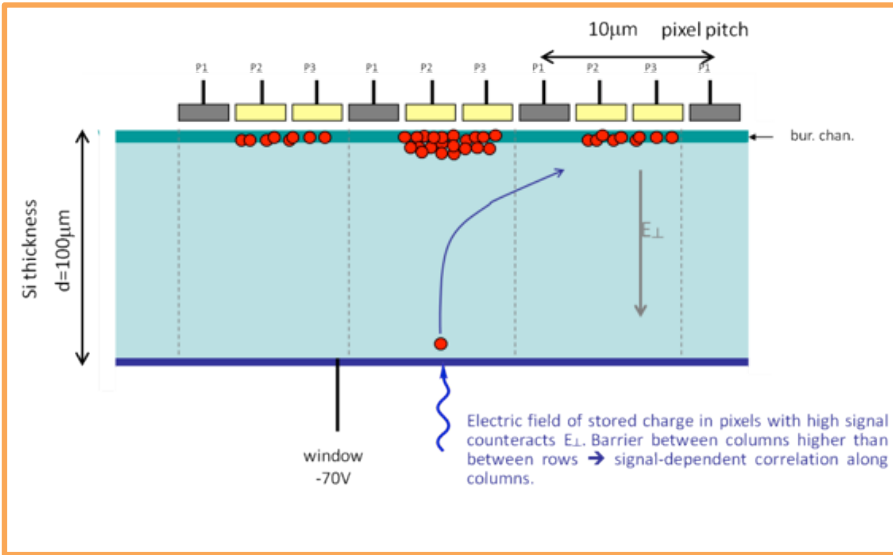
(c)



(d)

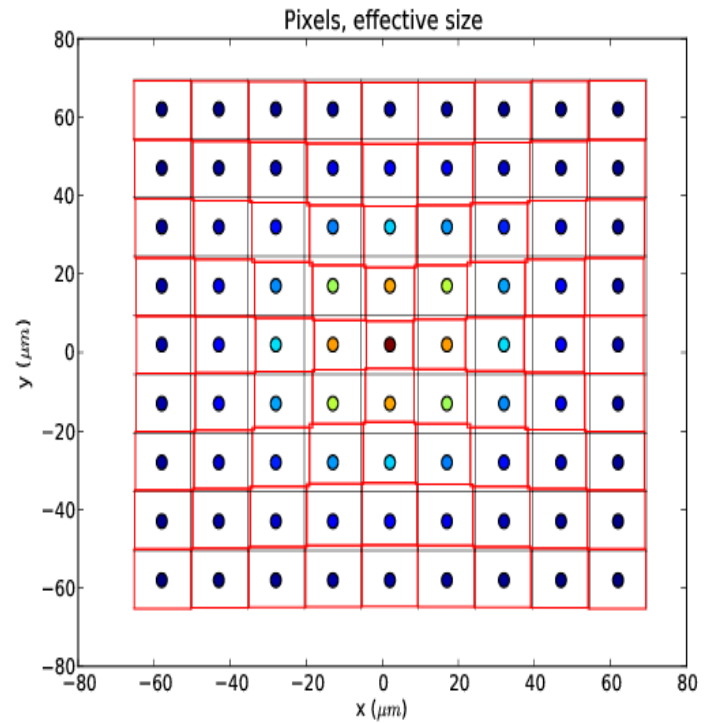


Major effort to study BF effect



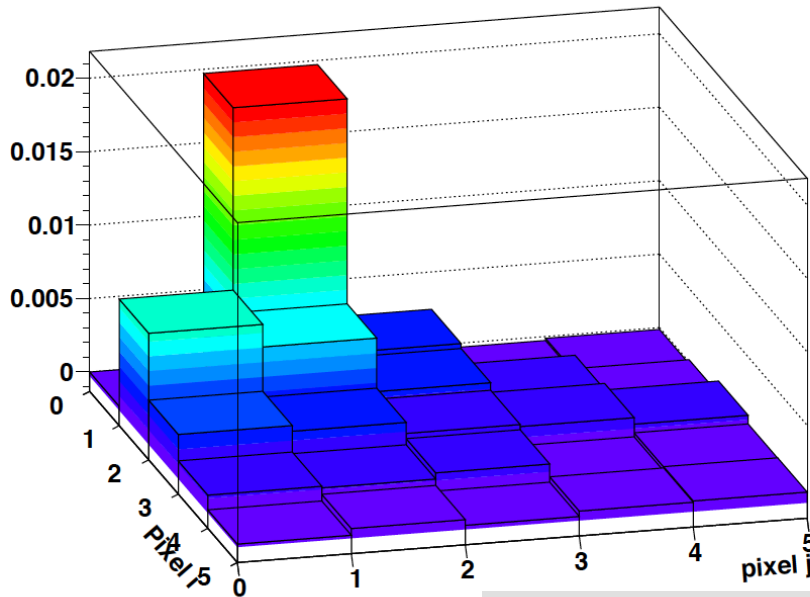
correlation CCD E2V @ 73ke

amp. #4
Entries 24



Correlations can be readily measured in flats and used to derive parameters of BF model for PSF

P Antilogus et al 2014 JINST 9 C03048



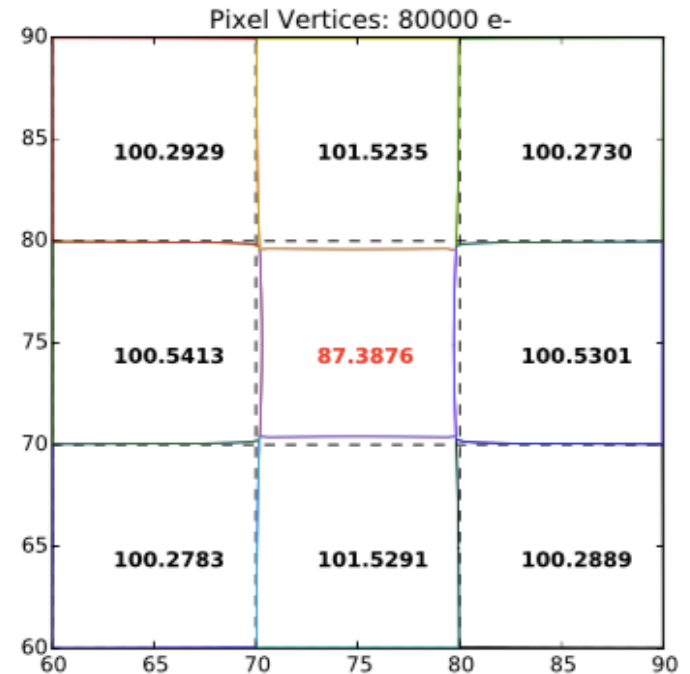
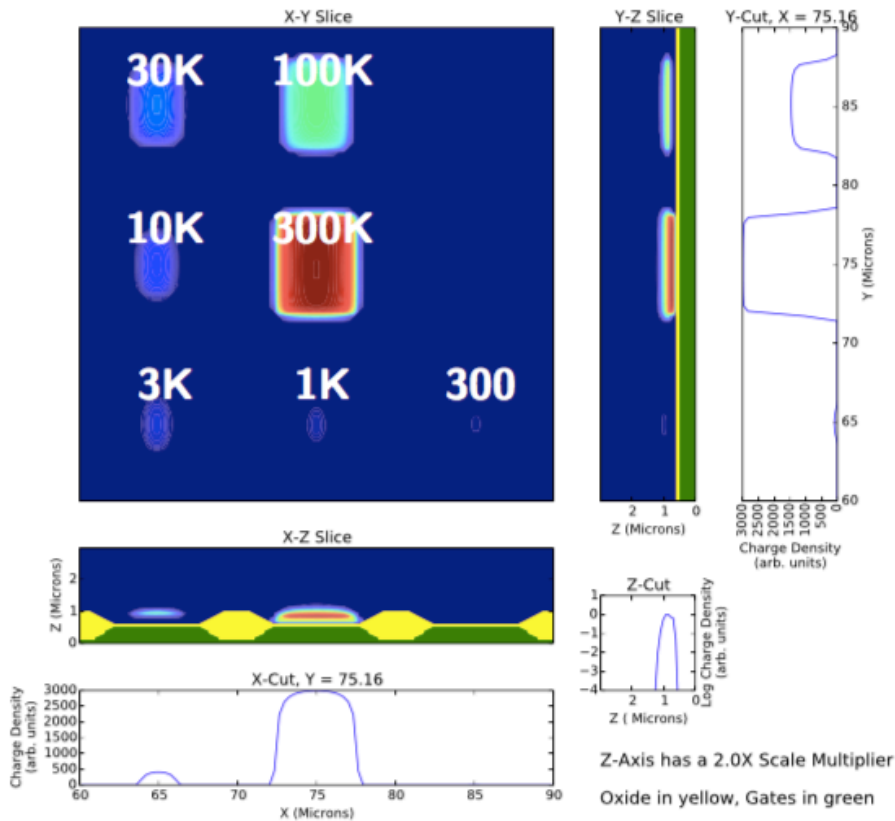
Pixel correlations

Poisson solver

Developed to model Brighter Fatter effect (Craig Lage, UC Davis)

C. Lage *et al* 2017 *JINST* **12** C03091

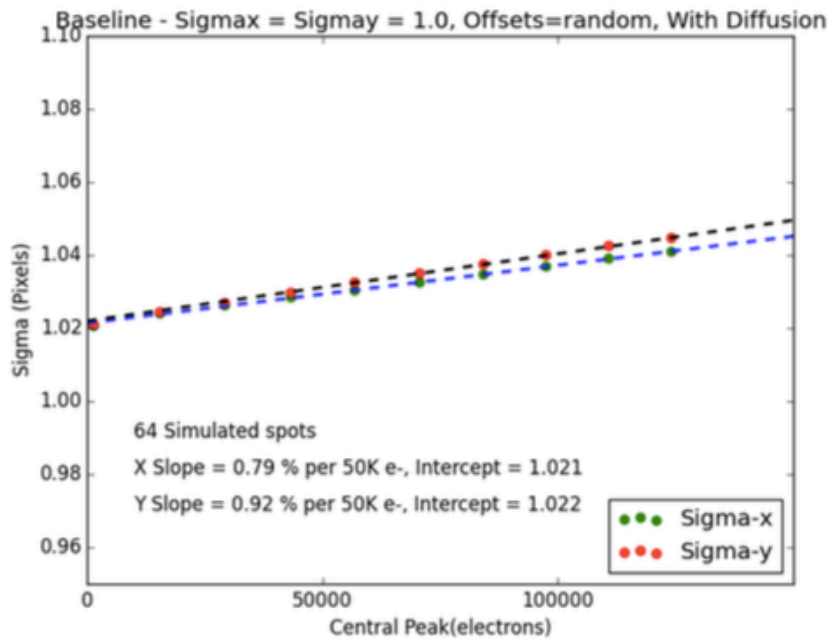
Electron Charge Distribution



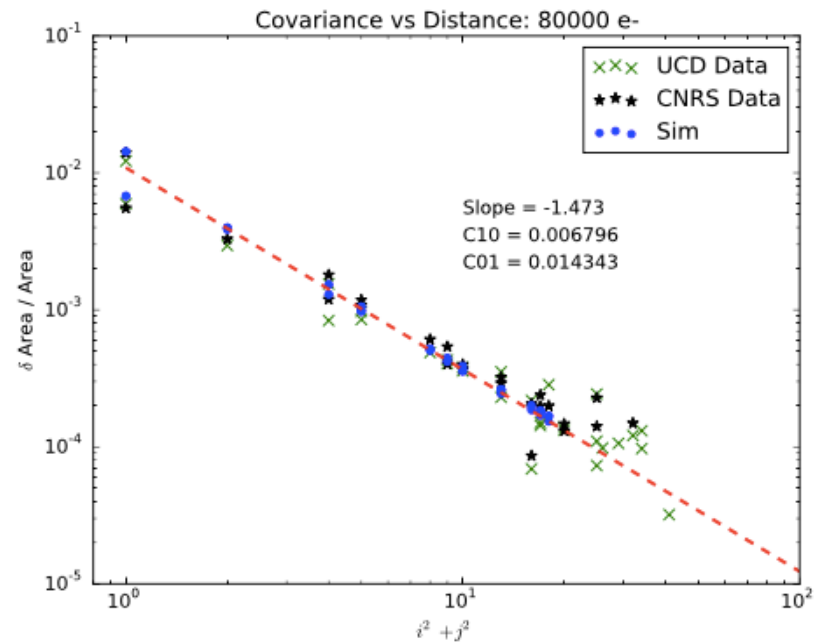
Success: same parameters describe PSF and correlations

Direct B-F - Measured:

$$X = 0.77 \pm 0.06; Y = 0.91 \pm 0.08$$

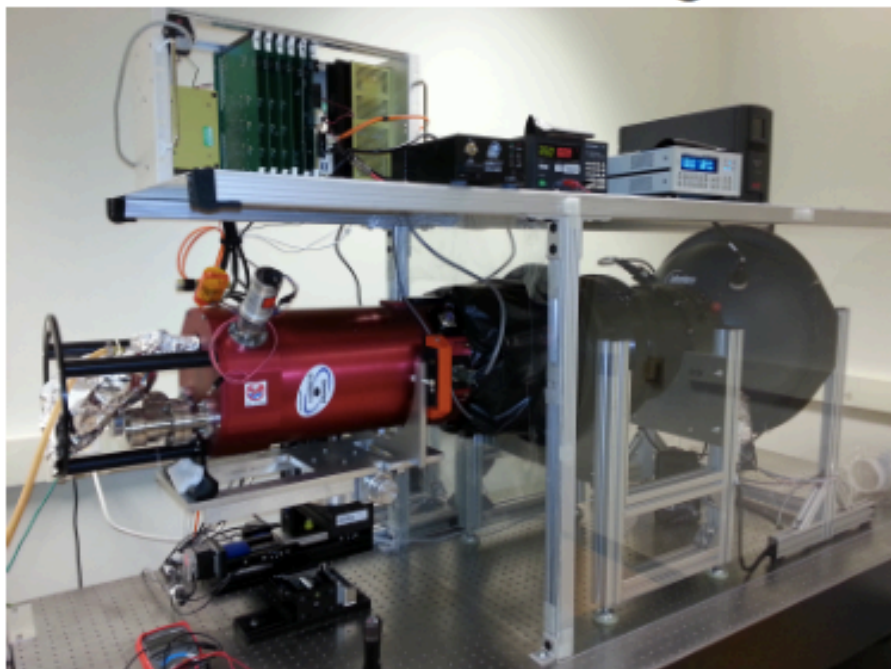


Pixel Correlations

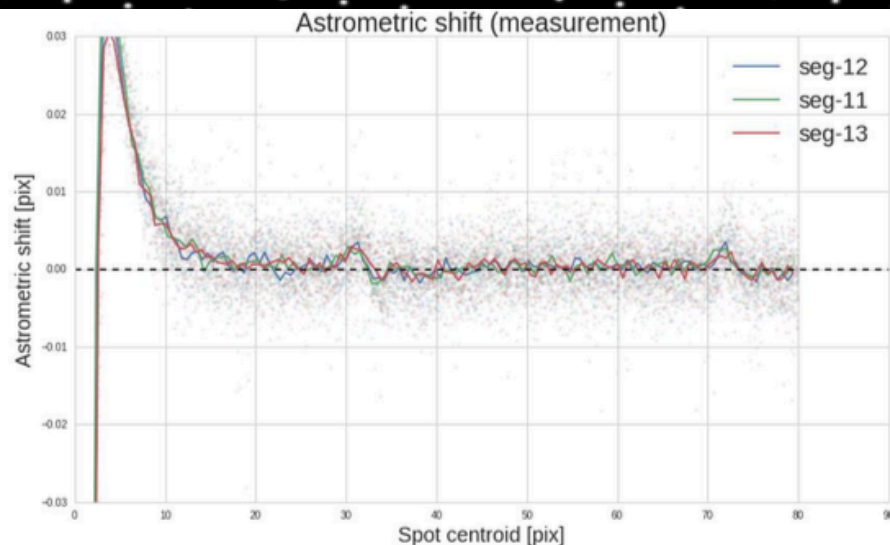
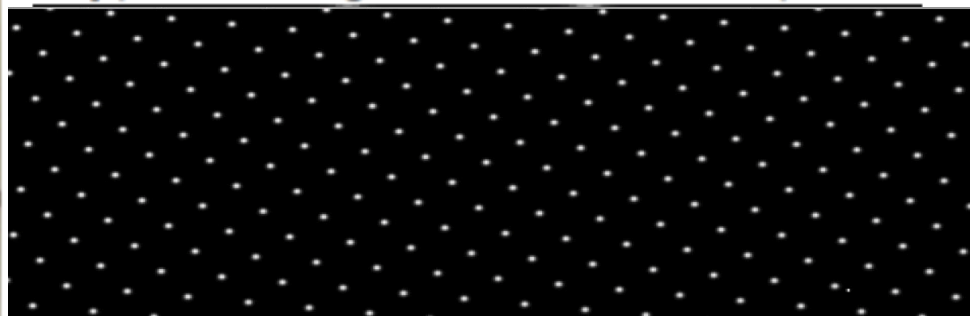


LSST Optical Simulator and Typical Spot Images

UC Davis 1:1 Re-Imager



Typical Image of 30 micron Spots:

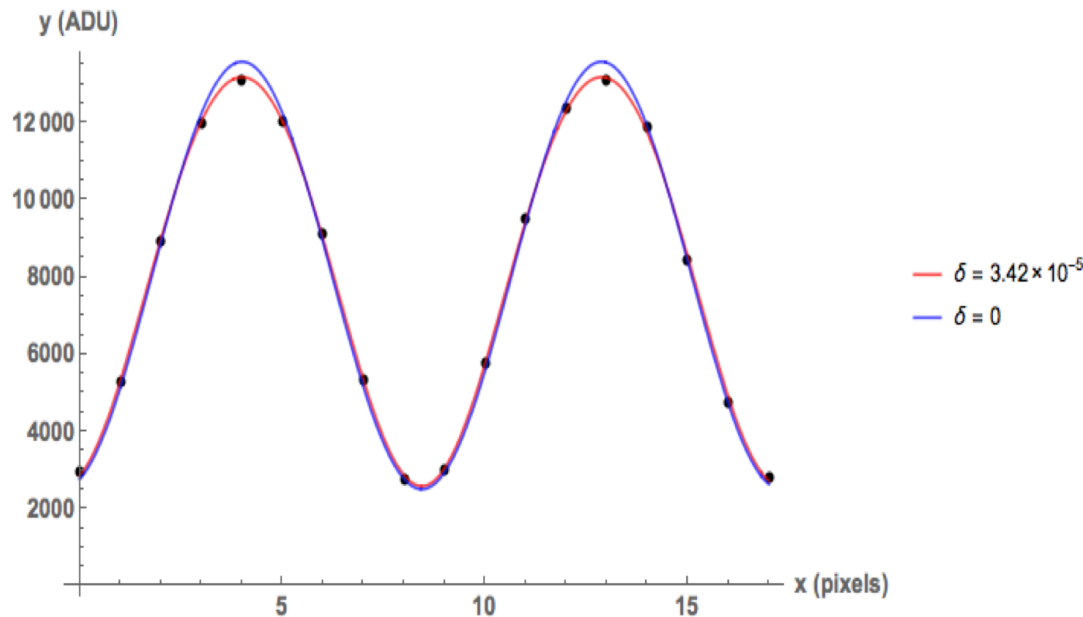
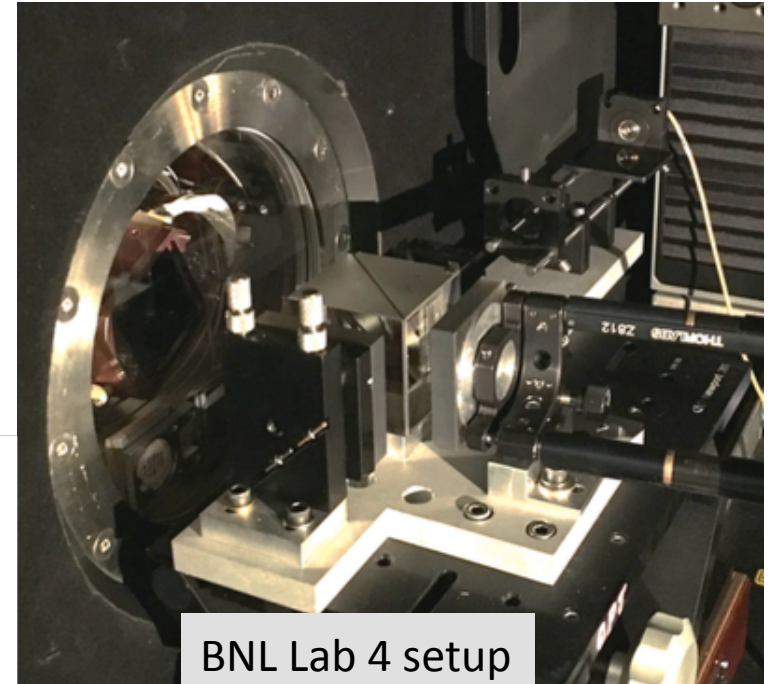


A. Bradshaw *et al* 2015 *JINST* **10** C04034

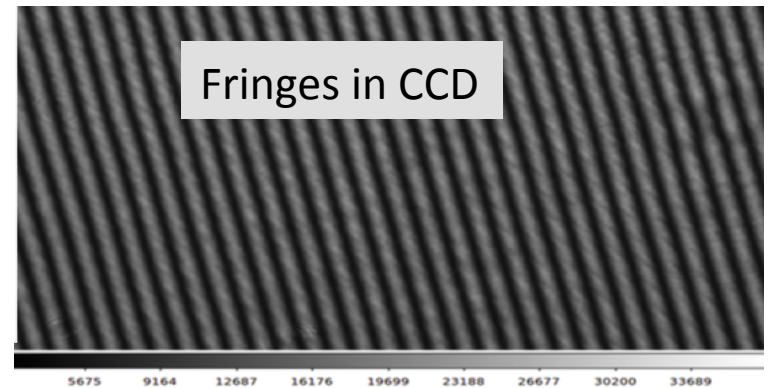
Tyson, et.al., "The LSST Beam Simulator", SPIE 9154-67 (2014), arXiv:1411.5667.

Fringe projector at BNL

- Michelson interferometer to generate sinusoidal variation of intensity
- Used to study Brighter Fatter effect



BF effect in fringes: data vs model

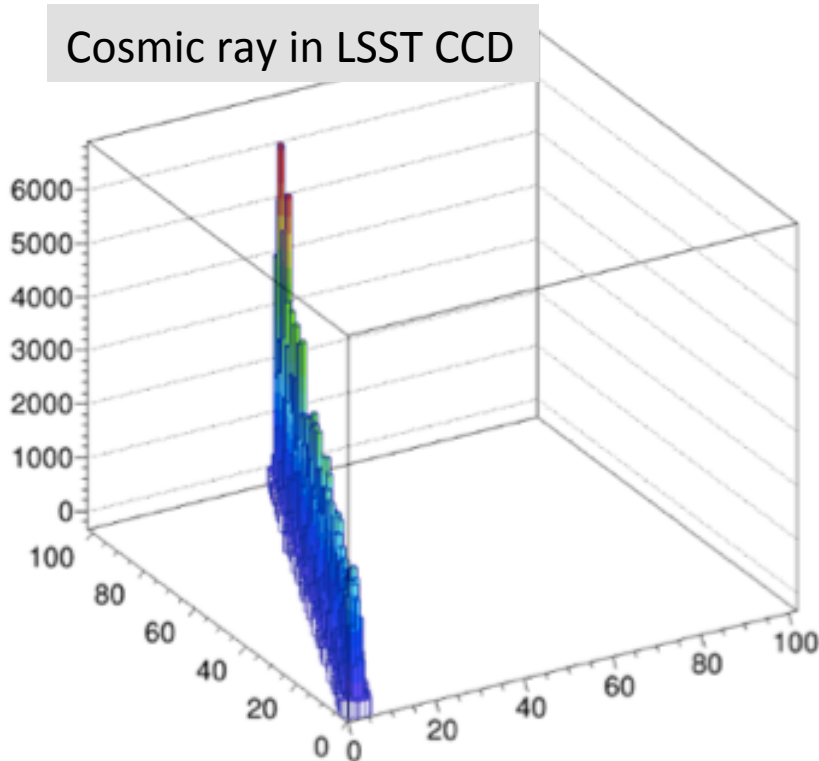


More CCD characterization

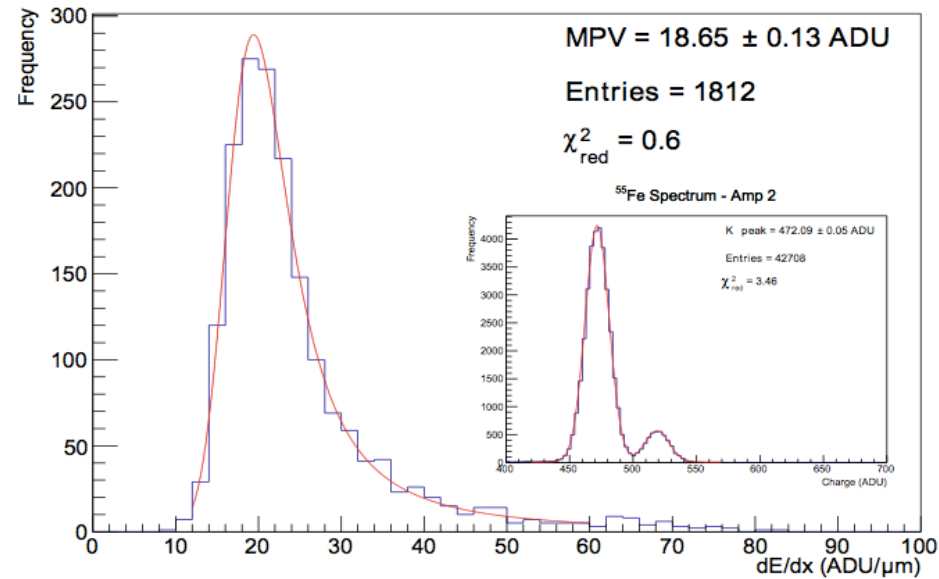
Used cosmics to measure gain and diffusion in LSST CCDs

dE/dx in CCD

Cosmic ray in LSST CCD



Muon energy deposition spectrum - Amp 2



CTE with Fe55

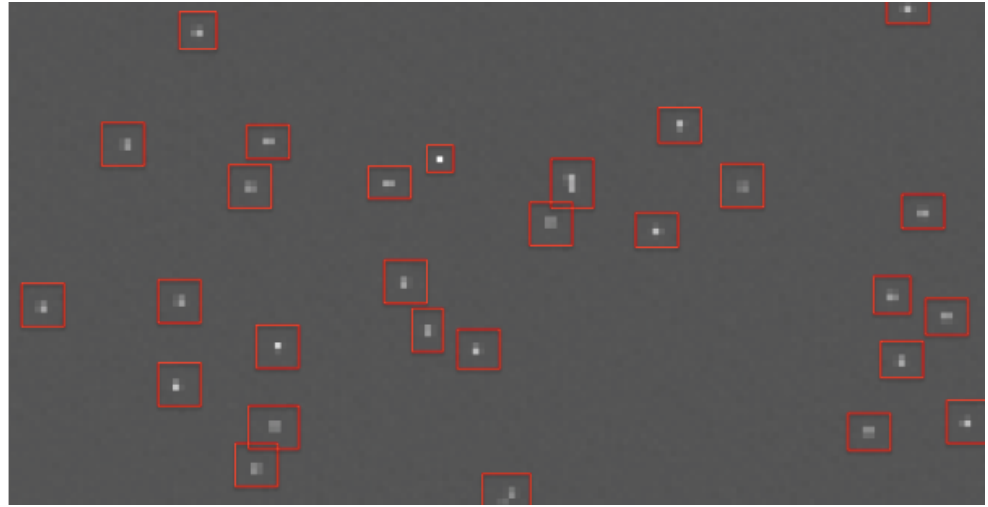
Using X-ray flats for CCD characterization

- Fe55 X-rays produce compact clouds of ~ 1600 electrons, $< 1 \mu\text{m}$
- Standard gain calibration technique for CCDs, used also for diffusion measurements
- Hit shape is symmetric but lateral electric fields in CCD can distort it (edges, tree rings etc)
- Uniform irradiation, not sensitive to the surface
 - 30 micron conversion depth
 - can extract astrometry and decouple it from photometry

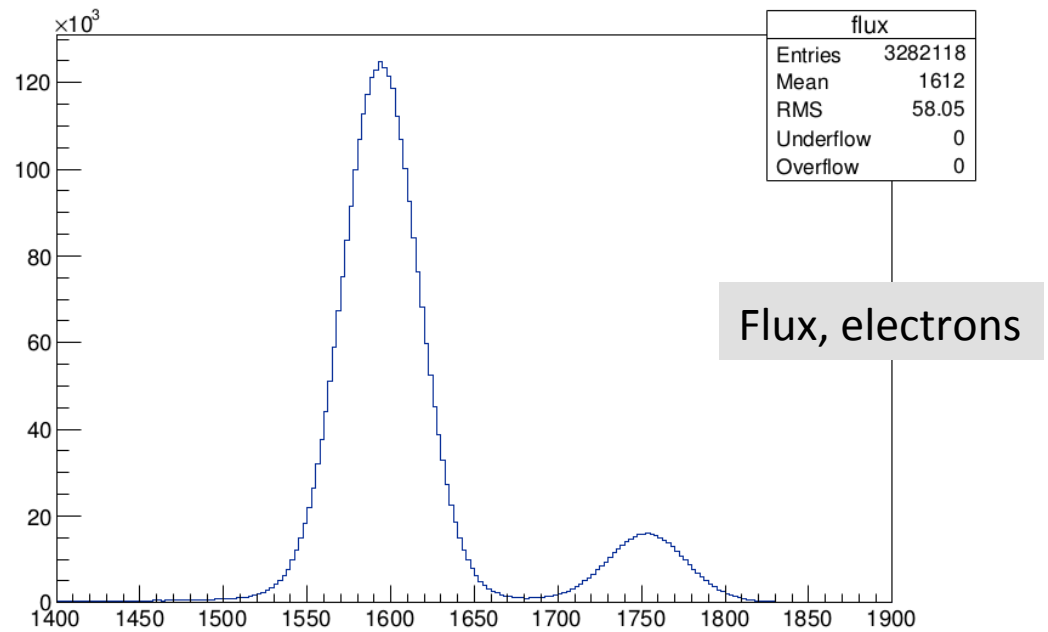
Easy to have good statistics, it's not too difficult to probe every pixel

→ X-ray flat fielding, dataset of 7M hits

CTE in Fe55 x-rays

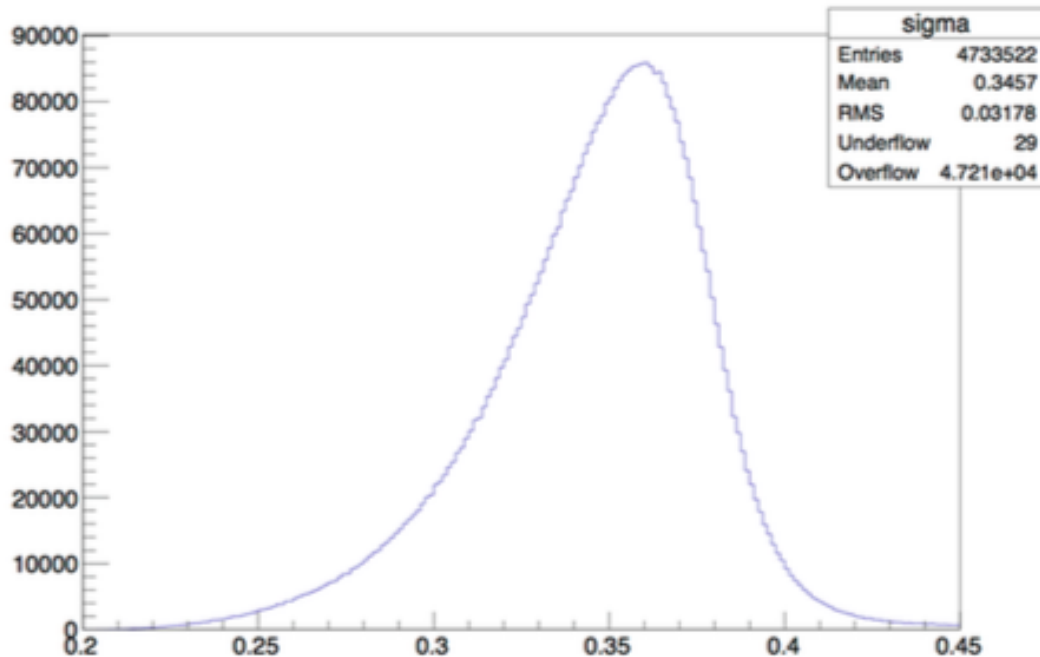


5.9 and 6.5 keV lines



Diffusion in CCD

3.6 micron in 100 micron thick sensor



Central pixel collects:

75 % for central hit
40% for side hit
25% for corner hit

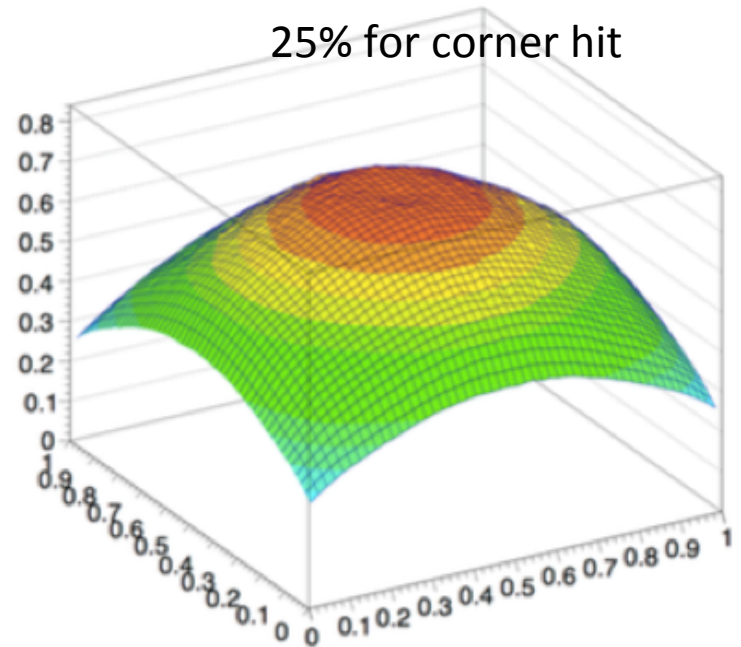


Figure 4 Distribution of the PSF size (sigma, in pixels) in the

2.0 μm
5.3 CTE MEASUREMENT TECHNIQUES

5.3.1 X-RAY TRANSFER

X-ray transfer, as discussed in Chapter 2, is the standard method in measuring absolute CTE performance.⁶ The technique is extremely valuable and is indispensable compared to other CTE measurement methods used which often give erroneous results. Figure 5.20 defines absolute CTE with the equation

$$CTE_X = 1 - \frac{S_D(e^-)}{X(e^-)N_P} \quad (5.19)$$

where CTE_X is the CTE as measured by x rays, $X(e^-)$ is the x-ray signal (e.g., 1620 e^- for an Fe^{55} x-ray source) and $S_D(e^-)$ is the average deferred charge after N_P pixel transfers (e^-).

Figure 5.21(a) shows a horizontal Fe^{55} x-ray transfer plot taken from a 520×520 CCD that exhibits a global CTE problem. The x-ray events interacting furthest from the on-chip amplifier experience the greatest charge loss. Note that the single-pixel-event line tilts "upward," whereas the deferred charge level for the first trailing pixel following the target pixel is seen as "downward" pattern

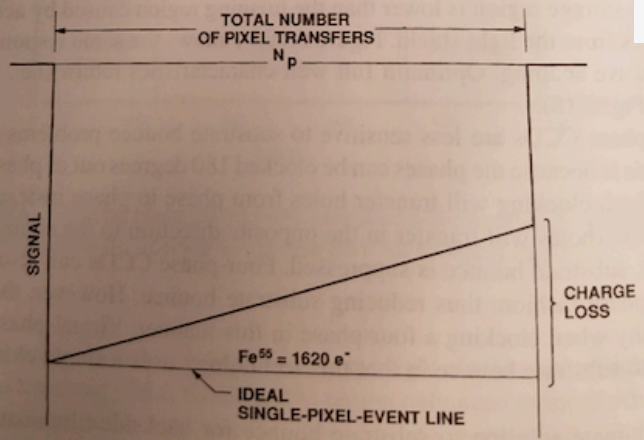


Figure 5.20 Absolute CTE definition using x-ray stimulation.

Janesick

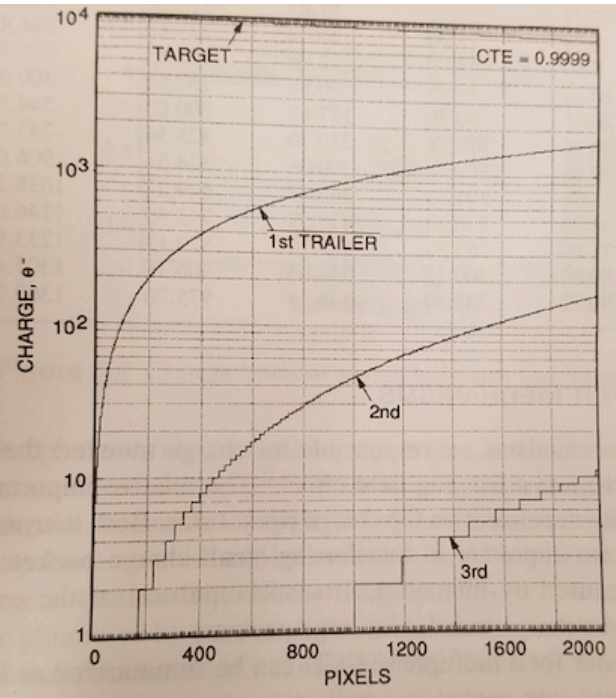
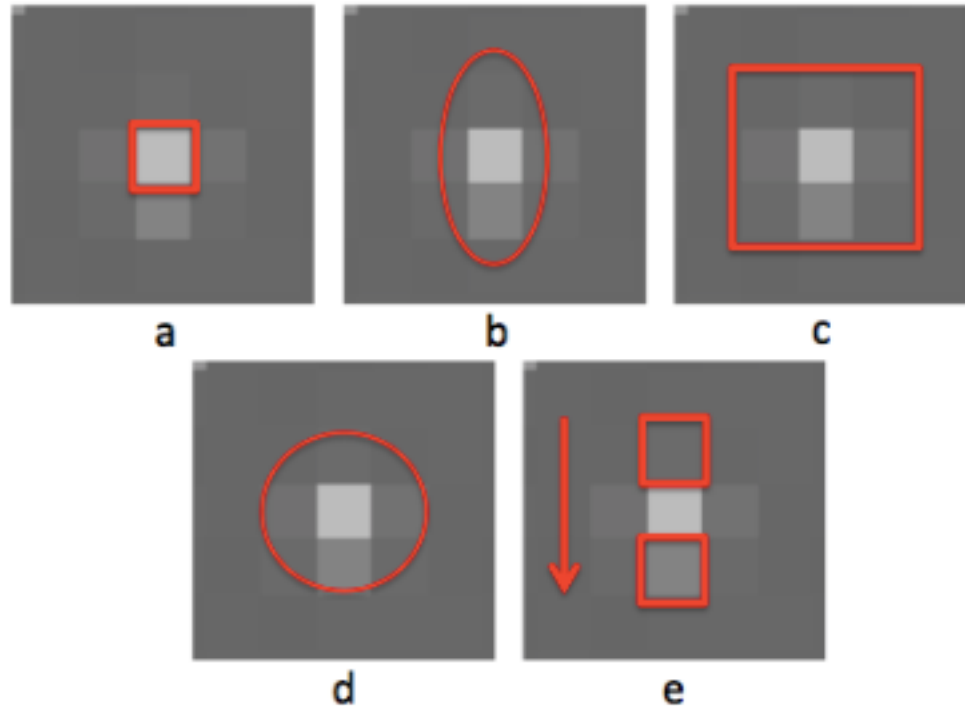


Figure 5.1(b) Theoretical 10,000- e^- point-source response assuming $CTE = 0.9999$.

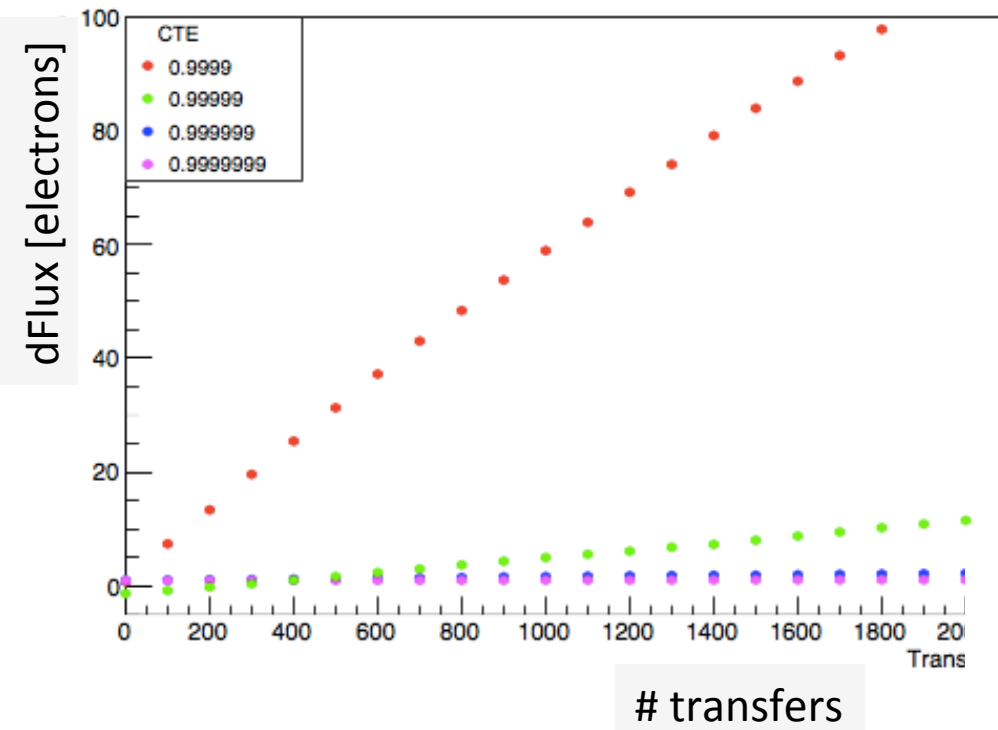
Problem: need to integrate several pixels to collect all charge → Start collecting deferred charge → Textbook formula is not applicable

CTE observables

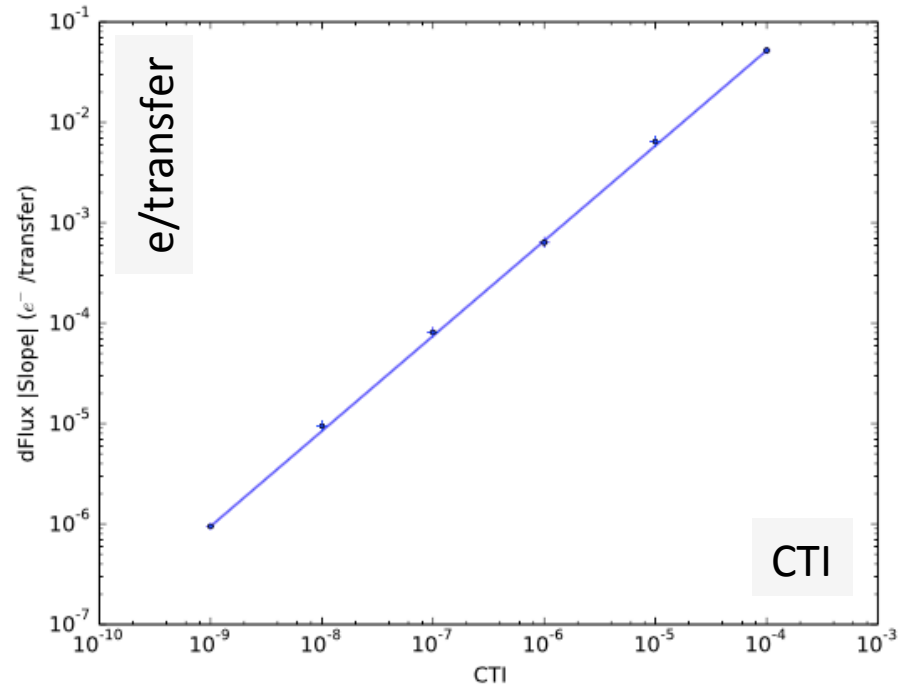


- MC simulations of Fe55 hits: assumes diffusion and random hit position
- Models deferred charge due to CTE
- Used this MC to predict behavior of observables as function of CTI

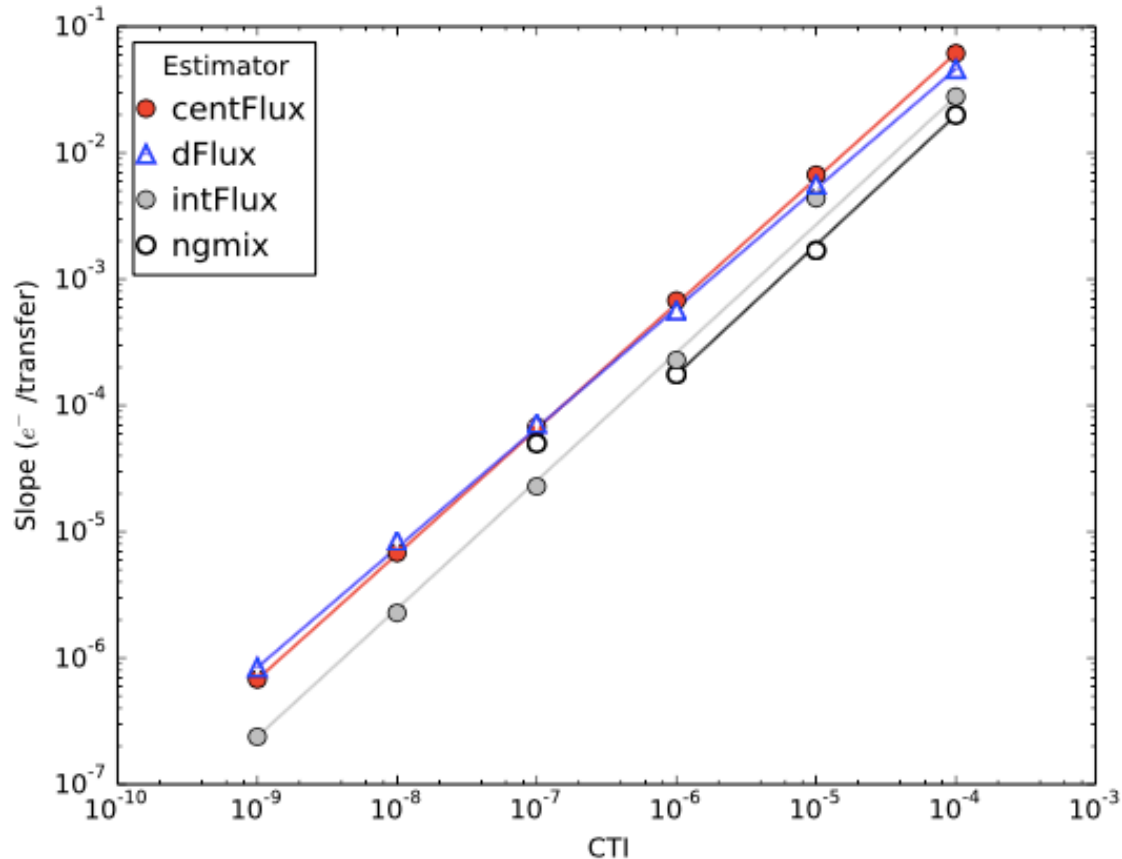
Differential Flux in MC



- Fit slopes for different CTI
- Plot slope [e⁻/transfer] vs CTI

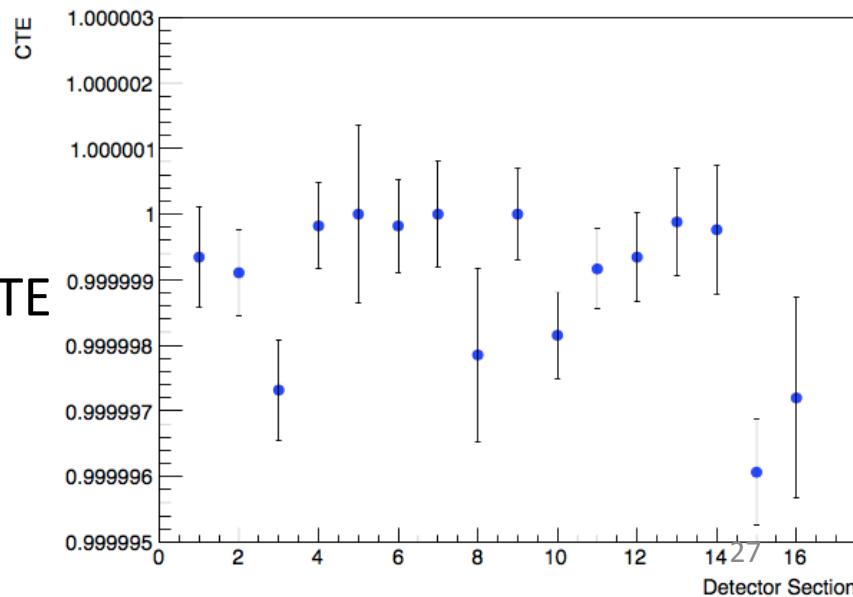
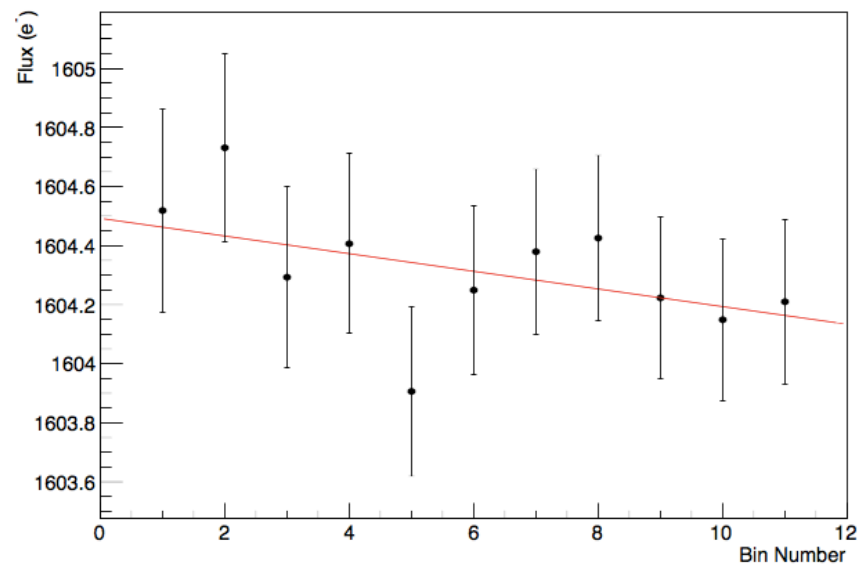
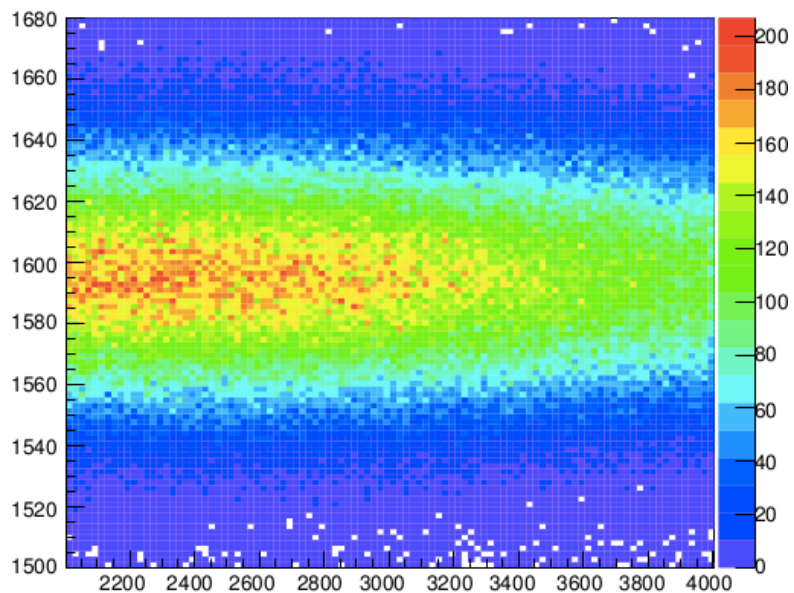


Slope vs CTI for all observables



Differential variables are slightly more sensitive

Example of parallel CTI in data



Flux [electrons] along columns

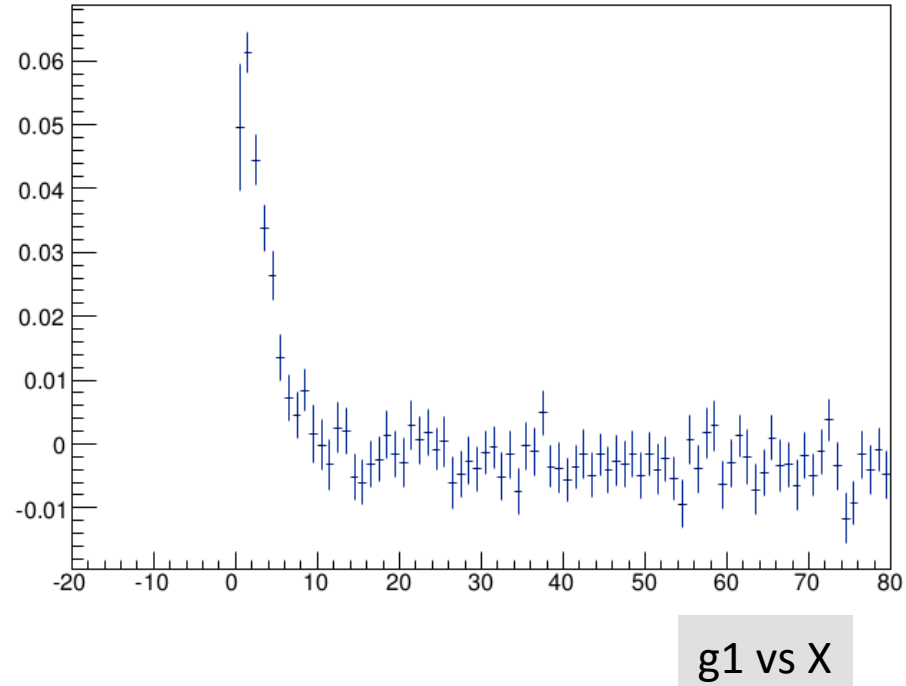
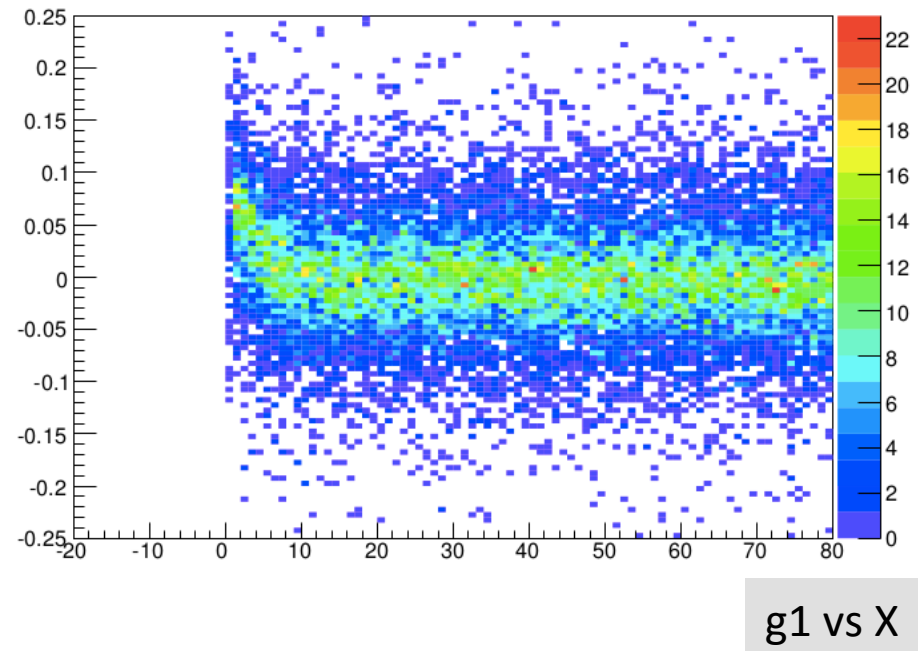
1. Measure slope
2. Apply MC results to convert slope to CTE

LSST spec for parallel CTE: 0.9999997

Fe55 hit shape

Edge Effect in Fe55 hit shape

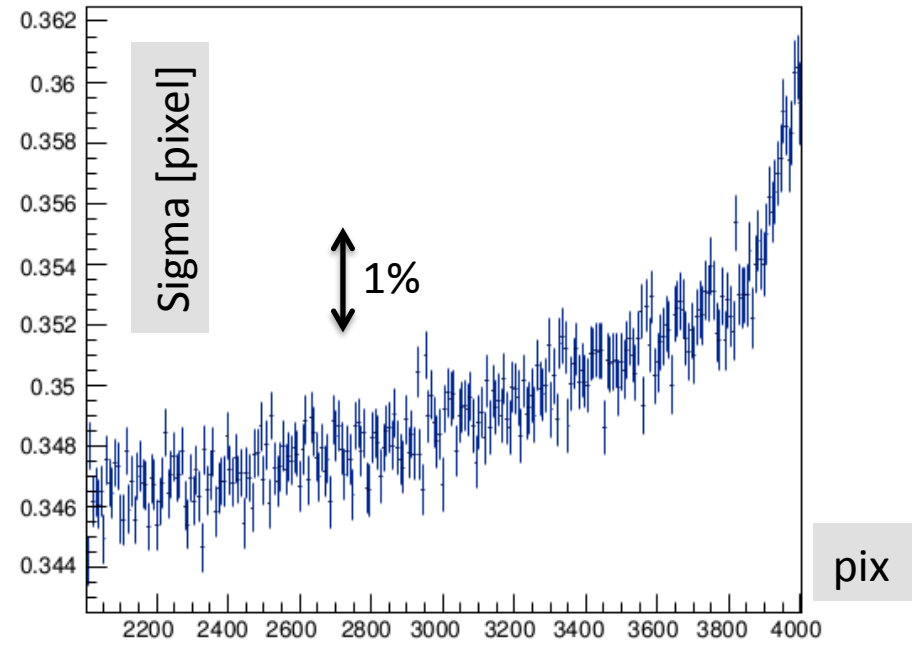
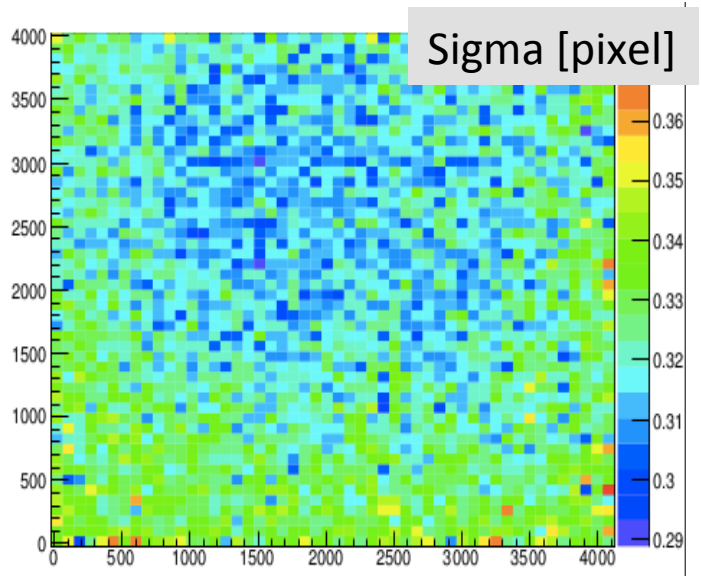
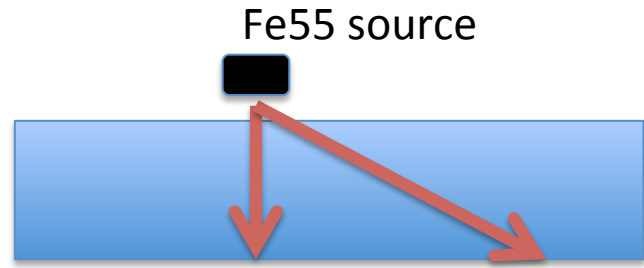
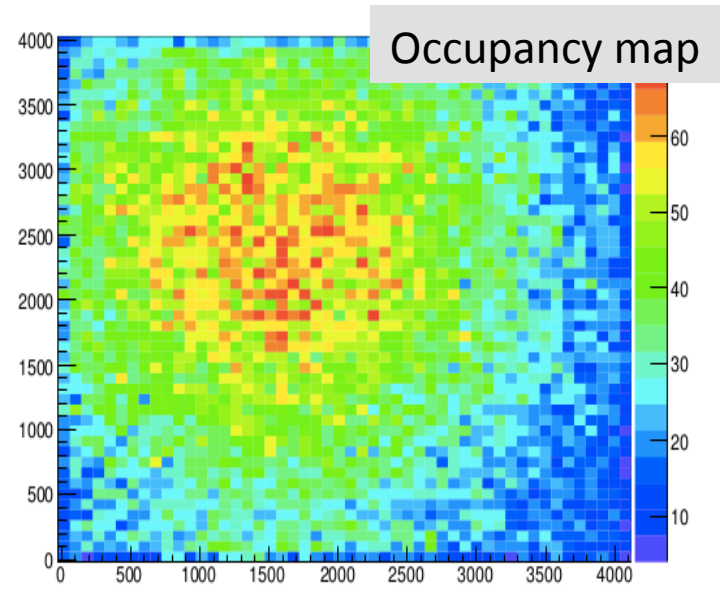
$$\text{Shear } g1 = (a-b)/(a+b)$$



$g1$ positive \rightarrow elongation along x , affects ~ 10 pixels

Sigma (x,y) Map

Center has smaller sigma since average depth of conversion is deeper there, hence less diffusion



Summary

Thick, fully depleted CCDs have non-trivial electrostatics → leads to distortions of pixel areas, corrections are important for science

LSST has a plan to study and remove these sensor signatures in the data and to model them correctly in science simulations

Fe55 flat fielding is a powerful tool to characterize CCDs

Fe55 can be used to measure CTE even for thick CCDs

Acknowledgements: Paul O'Connor, Ivan Kotov, Peter Takacs, Craig Lage, Andrew Bradshaw, Pierre Astier, Pierre Antilogus, Andy Rasmussen, Dan Weatherill, Merlin Fisher-Levine, Daniel Yates, Jason Brooks, Hyeyun Park, Homer Neal, Woody Gilbertson, Erin Sheldon